

# New magnetic field measurements of $\beta$ Cephei stars and slowly pulsating B stars

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**Abstract.** We present the results of the continuation of our magnetic survey with FORS 1 at the VLT of a sample of B-type stars consisting of confirmed or candidate  $\beta$  Cephei stars and Slowly Pulsating B stars. Roughly one third of the studied  $\beta$  Cephei stars have detected magnetic fields. The fraction of magnetic Slowly Pulsating B and candidate Slowly Pulsating B stars is found to be higher, up to 50%. We find that the domains of magnetic and non-magnetic pulsating stars in the H-R diagram largely overlap, and no clear picture emerges as to the possible evolution of the magnetic field across the main sequence.

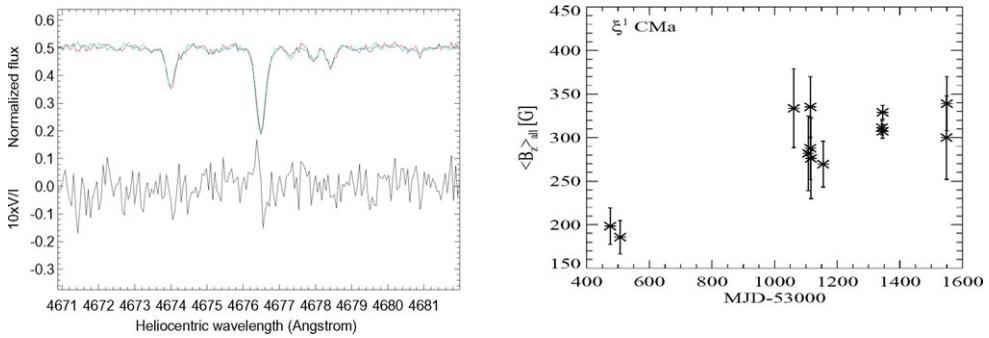
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## 1. Introduction

In our first publication on the magnetic survey of pulsating B-type stars (Hubrig *et al.* 2006), we announced detections of weak mean longitudinal magnetic fields of the order of a few hundred Gauss in 13 Slowly Pulsating B (hereafter SPB) stars and in the  $\beta$  Cephei star  $\xi^1$  CMa.  $\xi^1$  CMa showed a mean longitudinal magnetic field of the order of 300 G. Since the role of magnetic fields in modeling oscillations of B-type stars remained to be studied, we collected additional 98 magnetic field measurements in a sample of 60 pulsating and candidate pulsating stars.

## 2. Observations and analysis

The spectropolarimetric observations have been carried out in the years 2006 to 2008 at the European Southern Observatory with FORS 1 mounted on the 8-m Kueyen telescope of the VLT. For the major part of the observations we used the GRISM 600B in the wavelength range 3480–5890 Å to cover all hydrogen Balmer lines from H $\beta$  to the Balmer jump. One additional high resolution observation of  $\xi^1$  CMa has been obtained with the SOFIN echelle spectrograph installed at the 2.56 m Nordic Optical Telescope. The Stokes *I* and *V* spectra are presented in the left panel of Fig. 1. In the right panel of Fig. 1 we present magnetic field measurements of the same star using FORS 1 over the last 4.4 years. Among the SPBs and  $\beta$  Cephei stars, the detected magnetic fields are mainly of the order of 100–200 G.



**Figure 1.** Left panel: High-resolution ( $R \approx 30\,000$ ) SOFIN polarimetric spectrum of  $\xi^1$  CMa. A clear Zeeman feature is detected at the position of the unblended line O II  $\lambda 4676.2$ . Right panel: Magnetic field measurements with FORS 1 of  $\xi^1$  CMa over the last 4.4 years

### 3. Discussion

Roughly one third of the  $\beta$  Cephei stars have detected magnetic fields: Out of 13  $\beta$  Cephei stars studied to date with FORS 1, four stars possess weak magnetic fields, and out of the sample of six suspected  $\beta$  Cephei stars two show a weak magnetic field. The fraction of magnetic SPBs and candidate SPBs is found to be higher: roughly half of the 34 SPB stars have been found to be magnetic and among the 16 candidate SPBs, eight stars possess magnetic fields.

About a dozen pulsating stars discussed in our previous study Hubrig *et al.* (2006) have been subsequently observed by Silvester *et al.* (this proceedings) with ESPaDOnS at CFHT and NARVAL at TBL. The interesting result is that they could confirm our previous non-detections in eight  $\beta$  Cephei and SPB stars but failed to detect weak magnetic fields in the other five magnetic SPB stars selected from our study. On the other hand, as has been shown at the present meeting by Henrichs *et al.* (this proceedings) who confirmed our detection of a magnetic field in the B3V star 16 Peg, the integrated Zeeman features obtained from the application of a cross-correlation method, the Least-Squares Deconvolution, are extremely weak for  $\sim 100$  G fields. This means, that even using high-resolution spectrographs like ESPaDOnS or Narval, very high signal-to-noise observations are mandatory to be able to detect magnetic fields in hot stars where only a limited number of metal lines are suitable for the measurements.

In an attempt to understand why only a fraction of the pulsating stars exhibit magnetic fields, we studied the position of magnetic and non-magnetic pulsating stars in the H-R diagram. We find that their domains in the H-R diagram largely overlap, and no clear picture emerges as to the possible evolution of the magnetic field across the main sequence. It is possible that stronger fields tend to be found in stars with lower pulsating frequencies and smaller pulsating amplitudes. A somewhat similar trend is found if we consider a correlation between the field strength and the  $v \sin i$ -values, i.e. stronger magnetic fields tend to be found in more slowly rotating stars.

### References

Hubrig, S., Briquet, M., Schöller, M., *et al.* 2006, *MNRAS* 369, L61