

Numerical modeling of core-collapse supernovae and compact objects

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Abstract. Massive stars ($M \geq 10M_{\odot}$) end their lives with spectacular explosions due to gravitational collapse. The collapse turns the stars into compact objects such as neutron stars and black holes with the ejection of cosmic rays and heavy elements. Despite the importance of these astrophysical events, the mechanism of supernova explosions has been an unsolved issue in astrophysics. This is because clarification of the supernova dynamics requires the full knowledge of nuclear and neutrino physics at extreme conditions, and large-scale numerical simulations of neutrino radiation hydrodynamics in multi-dimensions. This article is a brief overview of the understanding (with difficulty) of the supernova mechanism through the recent advance of numerical modeling at supercomputing facilities. Numerical studies with the progress of nuclear physics are applied to follow the evolution of compact objects with neutrino emissions in order to reveal the birth of pulsars/black holes from the massive stars.

Keywords. stars: massive, stars: pulsars, stars: neutron, supernovae: general, methods: numerical, neutrinos, equation of state

1. Introduction

Supernova explosions are fascinating phenomena in the Universe (Bethe 1990). First of all, they have bright displays by gigantic explosions, some of which were recorded even in the old Chinese literature, as in the case of SN 1054 (the Crab nebula). They are named as *guest stars*, which is the same expression for the description of SN 1054 noted in the old Japanese literature. Secondly, they are associated with bursts of neutrinos, as in the case of SN 1987A (Hirata *et al.* 1987). In this event, neutrinos from the supernova were detected at terrestrial detectors. They revealed the importance of neutrinos in the supernova mechanism. This detection also led to the Nobel Prize in Physics in 2002 (Koshihara 1992, 2003). Most importantly, the supernova is the birthplace of compact objects: neutron stars and black holes. It is essential to understand the supernova mechanism to know the distribution of compact objects from massive stars, together with the information on mass and rotation.

Numerical modeling of core-collapse supernovae is the heart of the problem in order to grasp the timeline from massive stars to pulsars. Massive stars of $\sim 20M_{\odot}$ evolves for $\sim 10^7$ years with the formation of Fe core at the end of life. Gravitational collapse of the Fe core leads to core bounce, launch of shock wave and a resulting explosion in ~ 1 s. At center, a proto-neutron star is born and evolves thermally by emitting supernova neutrinos for ~ 20 s (e.g. Pons *et al.* 1999). It further cools down by radiation for $\sim 10^2$ - 10^5 years as a cold neutron star (e.g. Page *et al.* 2004). What one wants to know is the connection between the stellar evolution and the neutron star modeling. To bridge a gap between massive stars and pulsars, one needs to perform the supernova simulations, which are the most difficult part of the problem. This article describes ongoing challenges to clarify the supernova explosions with multi-physics at extreme conditions in 3 dimensions toward the goal to reveal the birthplace of neutron stars.

2. Challenges to understand the supernova mechanism

Despite the longstanding efforts over four decades, it is still not a simple task to reproduce supernova explosions in numerical modeling (Janka 2012, Kotake *et al.* 2012a). Among many issues of the supernova modeling, three major aspects of the challenges are: exotic conditions, counteracting $\sim 1\%$ effects and multi-dimensions. To tackle these difficult problems, it is necessary to perform large-scale simulations on the latest supercomputers.

To explain the current challenges in supernova physics, let us look into the details of supernova mechanism. Starting from the Fe core of the massive star, the gravitational collapse begins due to electron captures and photo-dissociations. The former reactions produce neutrinos, which escape freely at the beginning. These neutrinos are trapped inside the core at high densities during the collapse. Further collapse leads to the core bounce at just above the nuclear matter density, where the nuclear repulsive force is dominant. The shock wave is launched at the core bounce due to the stiffness of dense matter. If the shock wave propagates outwardly and reach the surface of the Fe core, this is the supernova explosion, which leaves the proto-neutron star and emits the supernova neutrinos. The proto-neutron star settles down as a cold neutron star afterward. The most difficult part among these stages is the way from the core bounce to the explosion with the three major challenges in the following subsections.

2.1. Nuclear physics at extreme conditions

One of the key ingredients in the supernova modeling is nuclear physics at extreme conditions. It is mandatory to provide the properties of hot-dense matter (the equation of state) and the rates of neutrino reactions in the supernova core for numerical simulations. The density inside the supernova core becomes very high, being dense more than that in nuclei. The matter becomes very neutron-rich as compared with the stable nuclei such as ^{56}Fe . The temperature may rise beyond $\sim 10^{11}$ K (~ 10 MeV) in the central core, which is different from the cold neutron star case. In addition, it is necessary to consider the neutrino reactions under the influence of the hot-dense matter. One needs energy- and angle-dependent neutrino reaction rates for emission, absorption, scattering and pair processes. Neutrino processes via all targets including nucleons, nuclei and leptons must be considered (e.g. Burrows *et al.* 2006).

As for the equation of state (EOS), it is a challenging task to provide the data set of EOS for wide range of density, temperature and composition for the supernova simulations. Although there are many results of the EOS for cold neutron stars, the set of EOS for supernovae have been limited before (Lattimer & Swesty 1991). There are continuous efforts to provide nuclear data for supernovae with the advance of experimental and theoretical nuclear physics since the 1990s. Experimental data of unstable nuclei became available at radioactive nuclear beam facilities along with the advance of nuclear many body theories (Serot & Walecka 1986, Brockmann & Machleidt 1990). Systematic measurements of the radii of neutron rich isotopes are valuable inputs to constrain the interaction in the nuclear models for supernova EOS (Shen *et al.* 1998a, 1998b, 2011).

The two EOS tables by Lattimer-Swesty and Shen *et al.* have been popular in many supernova simulation studies. There is recent progress of the EOS tables at low and high densities. One direction is extension to include exotic degrees of freedom such as hyperons and quarks at high densities (Ishizuka *et al.* 2008, Sagert *et al.* 2009, Nakazato *et al.* 2010a). The other direction is extension to consider the mixture of nuclei at low densities (Hempel *et al.* 2012). Recent observations of neutron star masses and radii are valuable to constrain the behavior of EOS at high densities. The observation of

the massive neutron star of $\sim 2M_{\odot}$ (Demorest *et al.* 2010) is strong constraint to select the EOS sets for supernova simulations. Determination of the neutron star radii is also influential (Steiner *et al.* 2010) and needs further careful analysis with model dependence (Suleimanov *et al.* 2011).

2.2. Neutrinos in supernova dynamics

An important key for the successful explosion is the counteracting 1% effects. Neutrinos are playing a critical role to determine the energy budget in the supernova dynamics. Since the Fe core of $\sim 10^3$ km collapses to the neutron star of ~ 10 km, there is gravitational energy release, which amounts to $\sim 10^{53}$ erg. The explosion energy including kinetic, internal and potential energies is $\sim 10^{51}$ erg from observations. Although the explosion energy sounds just like a fraction of the gravitational energy release, the most of the energy are actually carried away as supernova neutrinos. The total energy of the emitted neutrinos amounts to $\sim 10^{53}$ erg from the observation of the supernova neutrinos. Therefore, only 1% is used for the explosion out of the total release of the emitted neutrinos. To understand this delicate role of neutrinos, careful evaluation of the neutrino-matter interaction is essential in every stage of the supernova dynamics.

It is helpful to describe further some of the important stages as examples of the 1% effects or the issue of 10^{51} erg. At the early stage after the bounce, the initial shock wave stalls on the way because the energy is used up during the propagation. After the core bounce, the shock wave must propagate through the Fe core up to its surface against the falling material of outer layers. The energy loss occurs due to Fe dissociation during the propagation. Although the shock wave is launched up to ~ 150 km, it stalls without success in the spherical simulations. Looking at the energy budget, the initial energy of the shock wave amounts to several 10^{51} erg by the gravitational energy of inner core. The energy loss, on the other hand, due to the Fe dissociation amounts to minus several 10^{51} erg. Since this loss uses up the initial energy, the hydrodynamical explosion in ~ 100 ms is not possible.

Additional assistance to bring the energy of 10^{51} erg are necessary to revive the stalled shock wave. The neutrino heating mechanism behind the stalled shock wave is one of the plausible scenarios. At ~ 100 ms after the core bounce, the stalled shock wave is hovering around 200 km with accretion of outer layers (Figure 1, left). At center, the proto-neutron star is born already and emits abundant neutrinos from its hot and lepton rich core. A part of them hit the material just below the shock wave and are absorbed by nucleons (and nuclei). This contributes to the heating through the energy transfer from the neutrinos to the matter. This heating energy amounts to several 10^{51} ergs, which is comparable to the other factors described above. The size of this effect depends in detail on the shock wave dynamics, the energy and flux of neutrinos, the amount of the target material and duration time. In order to examine this process precisely, one has to solve the neutrino transfer together with the hydrodynamics. Modern numerical simulations must clarify whether the neutrino heating mechanism can revive the stalled shock wave as seen originally in the explosion by Bethe & Wilson (1985) via the delayed explosion mechanism.

2.3. Multi-dimensions with supercomputing technology

In addition to the role of neutrinos, novel features in multi-dimensional simulations are pivotal in the supernova mechanism. These are also challenging problems, which need the latest technology of supercomputers. When one breaks the symmetry from 1D (spherical) to 2D (axial) and 3D in the numerical simulations, there are always new findings on the explosion mechanism due to hydrodynamical instabilities such as convections and

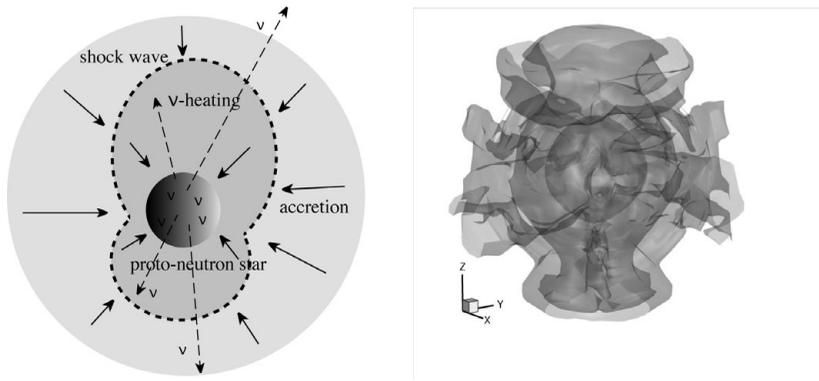


Figure 1. (Left) Schematic diagram of the deformed shock wave and neutrino heating. (Right) Neutrino distributions in a supernova core evaluated by neutrino transfer in 3D.

standing accretion shock instabilities (SASI) (Blondin *et al.* 2003). Moreover, the importance of asymmetry is supported by the observational facts on the asymmetric properties of supernova remnants. The main idea is the combination of the neutrino heating and the hydrodynamical instabilities. Because of the asymmetric launch of the shock wave (ex. up/down asymmetric), the shock wave along a certain direction reaches up to a larger distance than that realized in 1D. The neutrino heating below the deformed positions of the stalled shock wave helps to push the further propagation (Figure 1, left). The main gain is the sufficient amount of time for neutrino heating, since the advection time for the accretion of material becomes longer than the 1D case.

It is to be noted that solving neutrino radiation-hydrodynamics is necessary to clarify this problem. Radiation-hydrodynamics in multi-dimensions is generally a difficult problem in astrophysics and engineering. Along with the advance of supercomputers, there is step-by-step progress in the numerical treatment of neutrino-radiation hydrodynamics. In 1D (spherical symmetry), it is possible to perform the first-principle-type calculations (Rampp *et al.* 2000, Liebendörfer *et al.* 2001, Sumiyoshi *et al.* 2005). However, no explosion is found in the numerical simulations of various research groups for typical progenitors. In 2D, state-of-the-art simulations with detailed neutrino and nuclear physics produce a dozen of explosion cases (See Table 1 in Kotake 2012b). However, they have been done with approximate treatment of neutrino transfer. Although the neutrino heating mechanism is promising in most of the 2D simulations, there are also plausible scenarios under discussion. In order to determine the most dominant effect, further systematic studies in 2D are necessary.

Whether this kind of multi-dimensional effects remains important in 3D is the current topics in supernova studies. The advantage of 3D extensions over 2D models have been discussed in the systematic analysis by a simplified model (Nordhaus *et al.* 2010) but remains controversial with new results (Hanke *et al.* 2012). There will be also new instabilities, which is hidden in 2D but appear in 3D. With the rapid progress of supercomputing technology, it became recently to perform the first 3D simulation starting from the collapse of a massive star. Takiwaki *et al.* have studied the 3D core-collapse supernovae by the ray-by-ray approach with the IDSA approximation (Takiwaki *et al.* 2012). Explosions are found in 3D modeling with 3D motion of convection and deformed propagation of shock wave. However, it is still premature to conclude the intrinsic 3D effects since they have to perform more simulations with higher resolutions. One of the largest 3D simulations is running at the K-computer at Kobe, Japan, which has a peak speed of 10 PetaFLOPS (the world's 2nd fastest machine as of June 2012).

It is also necessary to replace approximate treatment of the neutrino transfer to the accurate one. It has been a formidable task to solve the neutrino-radiation transfer in 3D space since it is a 6D phase-space problem. Recently there is remarkable achievement to solve the 6D Boltzmann equation for neutrino distributions (Sumiyoshi & Yamada 2012) in supernova cores (Figure 1, right; see also Kotake *et al.* 2012a). Future ExaFLOPS supercomputers will enable full 3D simulations of neutrino-radiation hydrodynamics of core-collapse supernovae.

3. Toward the birth of pulsars

Until we get the solid mechanism for the successful explosion, one needs to assume explosions for astrophysical applications such as nucleosynthesis by parameterizing the shock wave launch. There are a number of applications to explore the neutron star properties after the supernova explosion.

By putting initial condition of a proto-neutron star born in supernovae, it is possible to follow the thermal evolution of the proto-neutron star through the cooling by emission of neutrinos (e.g. Pons *et al.* 1999). By predicting the neutrino spectra, average energies and luminosities as a function of time, one can explore the EOS influence as well as the dependence on the properties of proto-neutron stars. These are also important templates of supernova neutrino bursts for future detection of the next galactic supernova. The neutrino signals from the collapse of massive stars have been systematically studied including the case of black hole formations (e.g. Sumiyoshi *et al.* 2007). Their evaluation of spectra has been applied to predict the event numbers (Nakazato *et al.* 2010b) at the current and future facilities of neutrino detection (Kistler *et al.* 2011).

There have been studies of pulsar kicks and rotations through simulation of the parameterized explosion after the bounce (e.g. Blondin & Mezzacappa 2007). For example, Wongwathanarat *et al.* (2010) have studied the pulsar kicks and rotations by following the motion of accretion and ejecta, to predict the velocity of neutron stars through momentum conservation and gravitational pulls. They obtain the time evolution of pulsar kicks and angular momentum, and find a moderate number of ~ 300 km/s for kick velocities and ~ 600 ms for spin periods. In principle, it is necessary to connect this kind of outcome with the progenitor mass and rotations via full supernova simulations.

In summary, the study of core-collapse supernovae is fascinating with multi-physics from femto- to kilometer scales. It is a challenging problem regarding extreme conditions, counter-acting 1% effects of neutrinos and multi-dimensions. Understanding the mechanism of supernova explosions requires neutrino-radiation hydrodynamics in three dimensions. Implementation of microphysics on nuclear data into the numerical simulation is mandatory. With the progress of supercomputing resources, there are promising 2-D cases of explosions. The current challenge is the 3D numerical simulations to solve the neutrino-radiation hydrodynamics on the latest and future technology of supercomputers. These numerical challenges are indispensable for clarifying the essential cause of supernova dynamics, and to answer the birth of pulsars from the gravitational collapse of massive stars.

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References

- Bethe, H. A. & Wilson, J. R. 1985, *ApJ*, 295, 14
- Bethe, H. A. 1990, *Rev. Mod. Phys.*, 62, 801
- Blondin, J. M., Mezzacappa, A., & DeMarino, C. 2003, *ApJ*, 584, 971
- Blondin, J. M. & Mezzacappa, A. 2007, *Nature*, 445, 58
- Brockmann, R. & Machleidt, R. 1990, *Phys. Rev.*, C42, 1965
- Burrows, A., Reddy S., & Thompson, T. A. 2006, *Nucl. Phys.*, A777, 356
- Demorest, P. B., Pennucci, T., Ransom, S. M., Roberts, M. S. E., & Hessels, J. W. T. 2010, *Nature*, 467, 1081
- Hanke, F., Marek, A., Müller, B., & Janka, H.-T. 2012, *ApJ*, 755, 138
- Hempel, M., Fischer, T., Schaffner-Bielich, J., & Liebendörfer, M. 2012, *ApJ*, 748, 70
- Hirata, K., Kajita, T., Koshihara, M., Nakahata, M., & Oyama, Y. 1987, *Phys. Rev. Lett.*, 58, 1490
- Ishizuka, C., Ohnishi, A., Tsubakihara, K., Sumiyoshi, K., & Yamada, S. 2008, *J. Phys. G: Nucl. Part. Phys.*, 35, 085201
- Janka, H.-T. 2012, *Ann. Rev. Nucl. Part. Sci.* in press; arXiv:1206.2503
- Kistler, M. D., Yüksel, H., Ando, S., Beacom, J. F., & Suzuki, Y. 2011, *Phys. Rev.*, D83, 123008
- Koshihara, M. 1992, *Phys. Rep.*, 220, 229
- Koshihara, M. 2003, *Rev. Mod. Phys.*, 75, 1011
- Kotake, K., Sumiyoshi, K., Yamada, S., Takiwaki, T., Kuroda, T., Suwa, Y., & Nagakura, H. 2012a, *Prog. Theor. Exp. Phys.*, 2012, 01A301
- Kotake, K. 2012b, *Comptes Rendus Physique* in press; arXiv:1110.5107
- Lattimer, J. M. & Swesty, F. D. 1991, *Nucl. Phys.*, A535, 331
- Liebendörfer, M., Mezzacappa, A., Thielemann, F.-K., Messer, O. E., Hix, W. R., & Bruenn, S. W. 2001, *Phys. Rev.*, D63, 103004
- Nakazato, K., Sumiyoshi, K. & Yamada, S. 2010a, *ApJ*, 721, 1284
- Nakazato, K., Sumiyoshi, K., Suzuki H., & Yamada, S. 2010b, *Phys. Rev.*, D81, 083009
- Nordhaus, J., Burrows, A., Almgren, A., & Bell, J. 2010, *ApJ*, 720, 694
- Ott, C. D., Burrows, A., Dessart, L., & Livne, E. 2008, *ApJ*, 685, 1069
- Page, D., Lattimer, J. M., Prakash, M., & Steiner, A. W. 2004, *ApJS*, 155, 623
- Pons, J. A., Reddy, S., Prakash, M., Lattimer, J. M., & Miralles, J. A. 1999, *ApJ*, 513, 780
- Rampp, M. & Janka, H.-Th. 2000, *ApJ*, 539, L33
- Sagert, I., Fischer, T., Hempel, M., Pagliara, G., Schaffner-Bielich, J., Mezzacappa, A., Thielemann, F.-K. & Liebendörfer, M. 2009, *Phys. Rev. Lett.*, 102, 081101
- Serot, B. D. & Walecka, J. D. 1986, *Adv. Nucl. Phys.*, 16, 1
- Shen, H., Toki, H., Oyamatsu, K., & Sumiyoshi, K. 1998a, *Nucl. Phys.*, A637, 435
- Shen, H., Toki, H., Oyamatsu, K., & Sumiyoshi, K. 1998b, *Prog. Theor. Phys.*, 100, 1013
- Shen, H., Toki, H., Oyamatsu, K., & Sumiyoshi, K. 2011, *ApJS*, 197, 20
- Steiner, A. W., Lattimer, J. M. & Brown, E. F., 2010 *ApJ*, 722, 33
- Suleimanov, V., Poutanen, J., Revnivtsev, M., & Werner, K. 2011, *ApJ*, 742, 122
- Sumiyoshi, K., Yamada, S., Suzuki, H., Shen, H., Chiba, S., & Toki, H. 2005, *ApJ*, 629, 922
- Sumiyoshi, K., Yamada, S., & Suzuki, H. 2007, *ApJ*, 667, 382
- Sumiyoshi, K. & Yamada, S. 2012, *ApJS*, 199, 17
- Takiwaki, T., Kotake, K., & Suwa, Y. 2012, *ApJ*, 749, 98
- Wongwathanarat, A., Janka, H.-T., & Müller, E. 2010, *ApJ*, 725, L106